

Origin of the Elements or Elements of the Origin

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Near the **Beginning**

around

13.7 billion years ago,

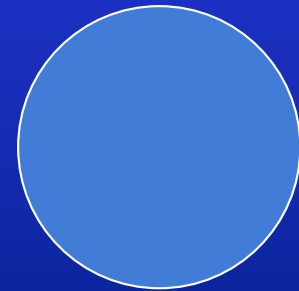
$t=10^{-43}$ seconds,

the 'Planck time'

$$\left(\sqrt{G \hbar / c^5} \right)$$

Quantum Fluctuations dominated Space-Time

The Universe might have
been
not much bigger
than
the Planck Length
 $r \approx 10^{-32} \text{ cm}$



*Lots of Energy
but no matter!*

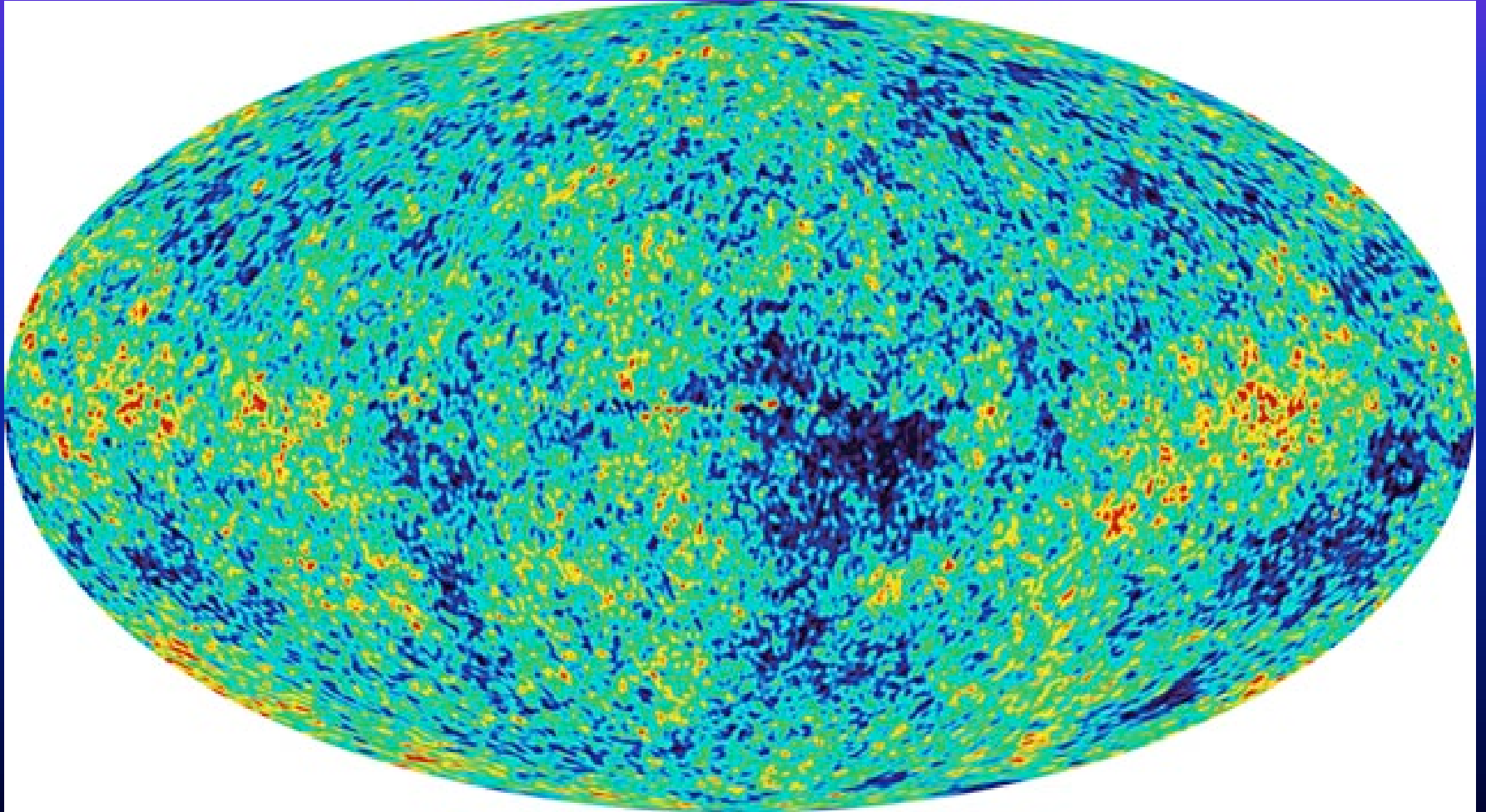
The energy could have
been contained in
a massless scalar field
and
the tension in space-time

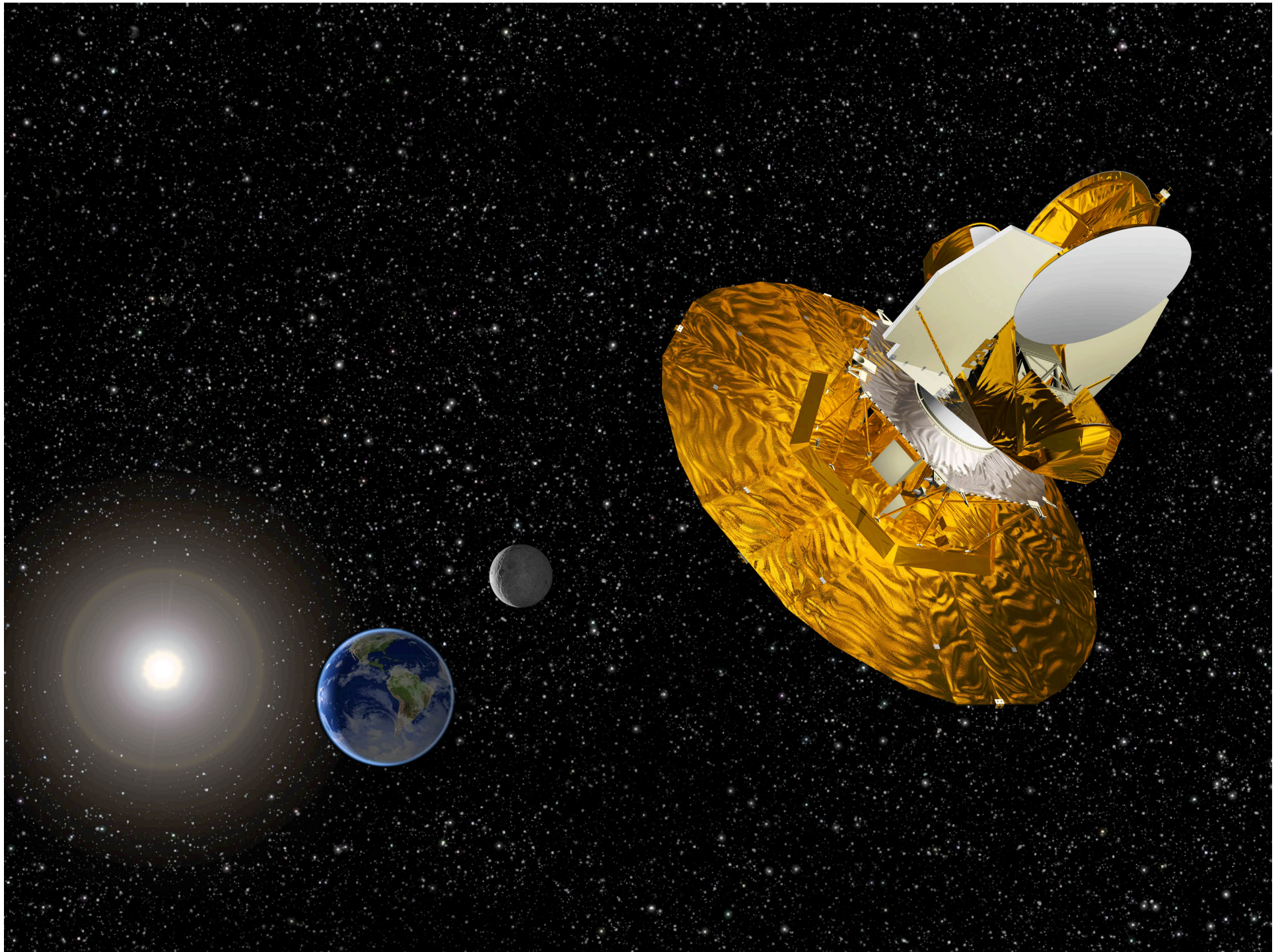
The scalar field
can cause a period
of exponential expansion,
'Inflation',
of the Universe
from $10^{-32}cm$ to $1cm$!

Inhomogeneities
'washed out'
to 1 part in 10^5
(superstring fluctuations also inflate)

Cosmic Background Explorer
COBE (1992) and
Wilkinson Microwave Anisotropy
Probe (2003)

WMAP of Universe





Scalar field undergoes
phase change:

Quantum Condensates

Super-heavy Bosons

Gluons, Photons, Weak Bosons,
Leptons, Quarks,
Baryons, Mesons

Structure Particles: Quarks and Leptons



	Q	M(GeV)
d	-1/3	0.006
u	2/3	0.006
s	-1/3	0.200
c	2/3	1.5
b	-1/3	5.1
t	2/3	178.

	Q	M(MeV)
ν_e	0	<0.003
e	1	.511
ν_μ	0	<0.19
μ	1	105.66
ν_τ	0	<18.2
τ	1	1777.

Interaction Generated by 'Gauge Bosons'

Particle	Symbol	Mass (GeV)
Super-heavy	X	$\sim 10^{15}$
Higgs	H^+, H^0	117.
Gluons	G	Zero
Weak Bosons	W^\pm, Z^0	80.42, 91.19
Photons	γ	$< 2 \times 10^{-16} \text{eV}$

By the elapse time of
 $t=10^{-6}$ seconds,
most baryons and antibaryons
annihilated.

1 in a billion left
as neutrons and protons

“Radiation Era”

Plasma of γ , ν , $\bar{\nu}$, e^- , e^+ , p , n

Temperature about 10^{12}K

Expansion continued more gently

$$R \propto t^{1/2}, T \propto t^{-1/2}$$

At $T \approx 10^{11} \text{K}$, (10^{-2} sec) neutrinos
decouple from matter

Below $T \approx 3 \times 10^{10} \text{K}$
deuterons
stick together to make Helium-4

At $T \approx 10^9 \text{K}$ (190 seconds),
positrons annihilate with
electrons to make photons



Neutrons taken into Helium
or decayed.

Left: γ , ν , $\bar{\nu}$, e^- , p , He^{++} (23% wt)
(He abundance sensitive to neutron lifetime,
 $\tau_{1/2} = 10.23$ minutes, and number of species of
neutrinos, $N=3$ from Z^0 decay!)

Beginning of 'Matter Era'

$t = 380,000$ years,

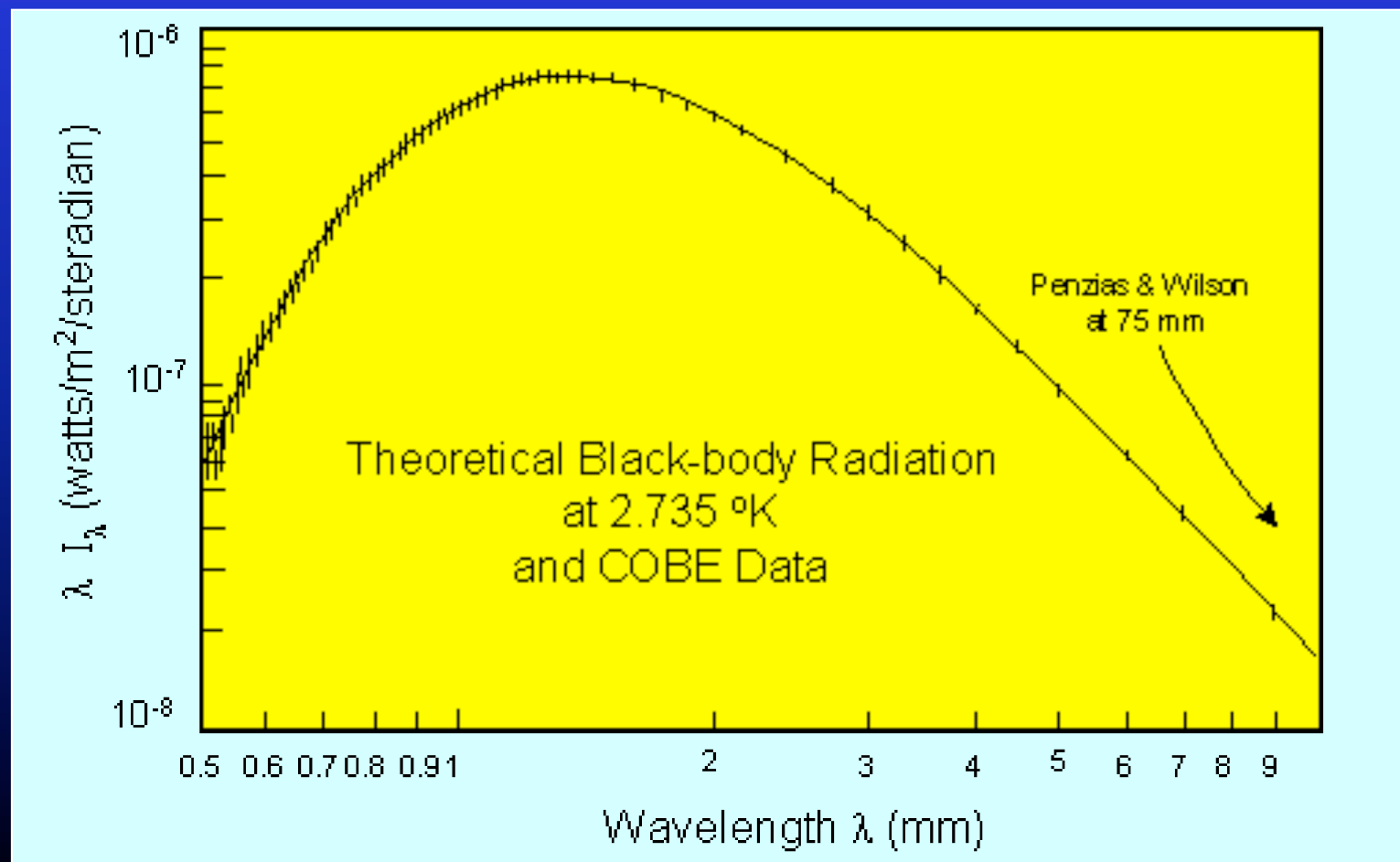
$T \approx 5000$ K

Matter 'decoupled' from radiation
when neutral atoms formed.

(Visible Matter 4%, Dark Matter 23%, Dark Energy 73%)

Background radiation predicted by
Gamow and Alpher (1948, GWU)

Residual Background found by Penzias and Wilson in 1965 COBE 2000 data:



'Stability Gaps' at A=5 and 8 stopped primordial nucleosynthesis

											¹² N -1.96	¹³ N -2.00	¹⁴ N 99.63	¹⁵ N 0.37
							⁸ C -20.7	⁹ C -0.90	¹⁰ C 1.29	¹¹ C 3.09	¹² C 98.9	¹³ C 1.1	¹⁴ C 11.26	¹⁵ C 0.389
						⁷ B -21.4	⁸ B -0.113	⁹ B -18.1	¹⁰ B 19.9	¹¹ B 80.1	¹² B 3.08	¹³ B 3.02	¹⁴ B 2.92	¹⁵ B 2.73
					⁶ Be -20.3	⁷ Be 6.66	⁸ Be -16.1	⁹ Be 100	¹⁰ Be 13.7	¹¹ Be 1.14	¹² Be -1.26		¹⁴ Be -2.40	
				⁵ Li -21.5	⁶ Li 7.5	⁷ Li 92.5	⁸ Li -0.076	⁹ Li -0.752		¹¹ Li -2.09				
		³ He .000138	⁴ He 99.9	⁵ He -21.1	⁶ He -0.093	⁷ He -20.5	⁸ He -0.924							
¹ H 99.985	² H 0.015	³ H 8.59												
¹ n 2.8				↑			↑							

Stable Element
Relative Abundance

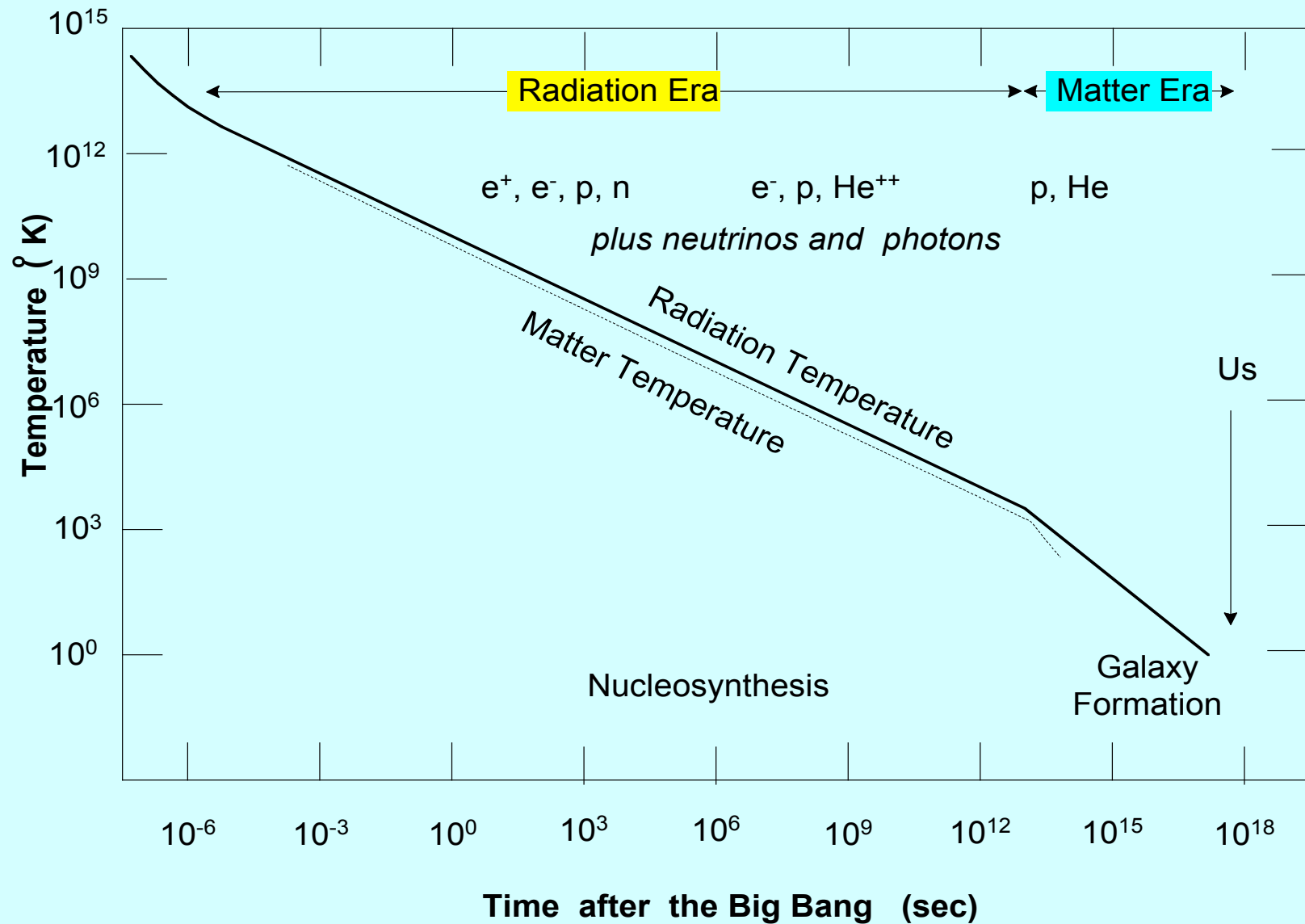
Unstable Element
LOG (Lifetime [sec])

Primordial Abundances of Light Elements

<i>Element</i>	<i>Mass Fraction, observed</i>	<i>Mass Fraction, theory</i>
^1H	0.75	0.76
^2He	$(2.5 \pm 1.5) \times 10^{-5}$	2.46×10^{-5}
^3He	$(4.2 \pm 2.8) \times 10^{-5}$	4.2×10^{-5}
^4He	0.23 ± 0.02	0.23
^6Li	$(320 \pm 200) \times 10^{-12}$	160×10^{-12}
^7Li	$(4800 \pm 3000) \times 10^{-12}$	270×10^{-12}

Data from spectroscopy of Population II (old) stars and gas

Time and Temperature



$t = 200$ million years

Galaxies and Stars form

Rate of heavier element production
strongly dependent
on nuclear stabilities, resonances,
temperature, and Coulomb barrier

Essence of Stellar Dynamics

Baryon, Lepton, Charge, & Energy Conservation

$$\frac{dL(r)}{dr} = 4\pi r^2 \rho(r) \varepsilon(\rho, T)$$

Local Mechanical Equilibrium

$$\frac{dp(r)}{dr} = -\rho(r) \frac{G}{r^2} \int_0^r \rho(r') 4\pi r'^2 dr'$$

Local Thermal Steady-State

Equation of State

$$p(r) = \frac{k}{\mu} \rho(r) T(r) + \frac{4\sigma}{3c} T(r)^4$$

Radiation flow

$$\frac{dT(r)}{dr} = - \frac{3\kappa\rho}{4\sigma T^3} \frac{L(r)}{4\pi r^2}$$

Convective flow

$$\frac{dT(r)}{dr} = \left(1 - \frac{c_v}{c_p}\right) \frac{T(r)}{p(r)} \frac{dp(r)}{dr}$$

Element Production in Stars

Reaction rate per unit number density:

$$\langle \sigma v \rangle = \sqrt{\frac{8}{\pi \mu (kT)^3}} \int_0^{\infty} \sigma(E) E e^{-\frac{E}{kT}} dE$$

Maximum rate:

$$E_0 = \left(\frac{\sqrt{2 \mu \pi} e^2 Z_1 Z_2 kT}{\hbar} \right)^{2/3}$$

Jumping the A=5 & 8 Barrier

Beryllium-8 resonance predicted by Fred Hoyle from
Carbon production via ${}^8\text{Be}(\alpha, \gamma){}^{12}\text{C}$

Resonance was later found.

Too short lived to help in early universe production.

Oxygen production: Subthreshold resonance in



Time Scales for Element Production in Stars

Hydrogen burning: 1 million to 20 million years

Helium burning: 100 thousand to 100 million years

Carbon burning: 500 to 1000 years

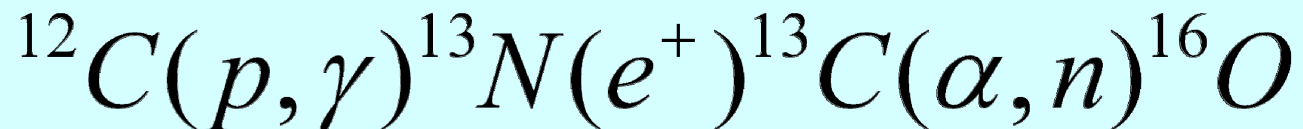
Neon burning: 1 year

Silicon burning: 1 day

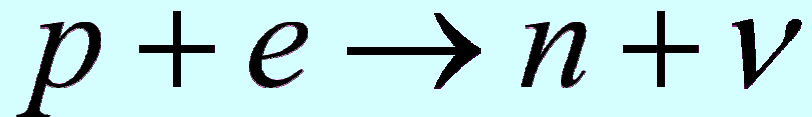
Iron made from excess neutrons
in reactions such as $^{22}\text{Ne}(\alpha, n)^{25}\text{Mg}$

Supernovae Element Production

Flux of excess neutrons, e.g.



and neutrinos



convert lighter elements to heavier ones in a
matter of seconds to days

Abundances of the Elements

